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LOFAR Properties of SILVERRUSH Ly α Emitter Candidates in the ELAIS-N1 Field

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ABSTRACT

Ly α emitters (LAEs) in the Epoch of Reionization (EoR) offer valuable probes of both early galaxy evolution and the process of reionization itself, however the exact evolution of their abundance and the nature of their emission remains an open question. We combine samples of 229 and 349 LAE candidates at $z = 5.7$ and $z = 6.6$ respectively from the SILVERRUSH narrowband survey with deep Low Frequency Array (LOFAR) radio continuum observations in the ELAIS-N1 field to search for radio galaxies in the EoR and study the low-frequency radio properties of $z \gtrsim 5.7$ LAE emitters. Our LOFAR observations reach an unprecedented noise level of $\sim 20 \mu\text{Jy beam}^{-1}$ at 150 MHz, and we detect 5 candidate LAEs at $> 5\sigma$ significance. Based on detailed spectral energy distribution (SED) modelling of independent multi-wavelength observations in the field, we conclude that these sources are likely [OII] emitters at $z = 1.47$, yielding no reliable $z \gtrsim 5.7$ radio galaxy candidates. We examine the 111 $z = 5.7$ and $z = 6.6$ LAE candidates from our panchromatic photometry catalogue that are undetected by LOFAR, finding contamination rates of 81-92% for the $z = 5.7$ and $z = 6.6$ subset of the LAE candidate samples. This subsets of the full sample is biased towards brighter magnitudes and redder near-infrared colours. The contamination rates of the full sample will therefore likely be lower than the reported values. Contamination of these optically bright LAE samples by likely [OII] emitters is lowered significantly through constraints on the near-infrared colours, highlighting the need for infrared observations to robustly identify bright LAEs in narrowband surveys. Finally, stacking of radio continuum observations for the robust LAE samples yields a 2σ upper limit on radio luminosity of 8.2×10^{23} and $8.7 \times 10^{23} \text{ W Hz}^{-1}$ at $z = 5.7/6.6$ respectively, corresponding to limits on their median star-formation rates of < 53 and $< 56 \text{ M}_{\odot} \text{ yr}^{-1}$.

Key words. Radio continuum: galaxies – galaxies: active – galaxies: high-redshift

1. Introduction

Luminous Lyman alpha emitters (LAEs) at high redshift offer valuable probes of galaxy evolution and cosmology in the early Universe. The existence of young galaxies in the early phase of their evolution with strong Ly α emission (1216 Å) was first hypothesised by Partridge & Peebles (1967) and first observed by Hu & McMahon (1996) and Pascarella et al. (1996). Dedicated narrowband, spectroscopic, and integral field unit surveys have since discovered thousands of LAEs at redshifts $2 < z < 7$, opening up a new way of studying the high redshift Universe (e.g. Cowie & Hu 1998; Rhoads et al. 2000; Ouchi et al. 2003; Ciarullo et al. 2011; Shibuya et al. 2012; Konno et al. 2014; Sobral et al. 2018b). Previous studies have identified two populations of LAEs: blue and faint LAEs with low mass and metallicity (e.g. Bacon et al. 2015; Sobral et al. 2015b, 2019; Nakajima et al.

2016; Ono et al. 2010) and red, massive and luminous LAEs (e.g. Chapman et al. 2005; Sandberg et al. 2015; Matthee et al. 2016) often observed to host an active galactic nucleus (AGN) (e.g. Ouchi et al. 2008; Konno et al. 2016; Sobral et al. 2017).

LAEs have been widely used to study luminosity functions and clustering properties of galaxies in the early Universe (e.g. Ouchi et al. 2003; Shimasaku et al. 2004; Ouchi et al. 2010; Kusakabe et al. 2018; Khostovan et al. 2019). Recent studies of Khostovan et al. (2019), for example, suggest that LAEs are progenitors of a wide range of galaxy types, where the brightest LAEs are located in the most massive halos and are highly clustered. LAEs that have been detected at radio wavelengths have been found to have steep radio spectral indices (Calhau et al. 2020).

Furthermore, studies have shown that the AGN fraction rises with Ly α luminosity, leading to the AGN dominating (over the

star-forming population) at the bright end of the $\text{Ly}\alpha$ luminosity function (e.g. Sobral et al. 2018a,b; Matthee et al. 2017; Wold et al. 2014, 2017; Calhau et al. 2020). As suggested by Sobral et al. (2018a) and demonstrated by Calhau et al. (2020), the AGN fraction of LAEs declines towards higher redshifts at fixed $\text{Ly}\alpha$ luminosity. However, these conclusions are limited by the number of detected LAE at high redshift and larger samples are required to confirm this claim. Constraining the AGN fraction in the $\text{Ly}\alpha$ luminosity function is vital for our understanding of early super-massive black hole (SMBH) formation (e.g. Calhau et al. 2020), but also critical for understanding the source of ionising photons of LAEs if they are to be used as probes of reionization (e.g. Santos et al. 2016; Matthee et al. 2015).

The LOFAR Two Metre Sky Survey (LoTSS; Shimwell et al. 2017, 2019) is entering a new regime of deep, low frequency surveys by pushing noise levels to below $100 \mu\text{Jy beam}^{-1}$ at 150 MHz across the entire Northern sky, and complementing this with targeted deeper observations in the degree-scale northern deep fields. The LoTSS deep fields first data release (Tasse et al. 2020; Sabater et al. 2020; Kondapally et al. 2020; Duncan et al. 2020a) covers a total area of 25 deg^2 in the Lockman Hole, ELAIS-N1 (hereafter EN1) and Boötes fields, at $6''$ resolution, reaching an RMS depth of $S_{150\text{MHz}} \sim 20 \mu\text{Jy beam}^{-1}$ in the deepest field EN1.

One of the LoTSS target fields, EN1 (RA=242.75, Dec=54.95 degrees), has also been the target field for a dedicated LAE survey named *Systematic Identification of LAEs for Visible Exploration and Reionization Research Using Subaru HSC* (SILVERRUSH; Ouchi et al. 2018). The SILVERRUSH program makes use of the narrowband observations of the Hyper Suprime-Cam Subaru Strategic Program (HSC-SSP; Aihara et al. 2018a). Three major scientific goals of the program are studying LAE properties at high redshift, using LAE to probe the low-mass young galaxy population, and studying the $\text{Ly}\alpha$ luminosity function towards the Epoch of Reionization (EoR) (Ouchi et al. 2018). The SILVERRUSH program identified ~ 2000 LAE candidates at $z = 5.7$ and $z = 6.6$ across 4 different fields, of which ~ 600 are located in EN1 (Shibuya et al. 2018b). The overlapping LoTSS and SILVERRUSH observations of EN1 open up a new opportunity to study the as yet unexplored low-frequency (150 MHz) radio properties of LAEs. Furthermore, the identification of LAE radio AGN would allow us to study the AGN fraction as a function of $\text{Ly}\alpha$ luminosity and follow up on the claim that the LAE AGN fraction declines towards higher redshift. Finally, according to model predictions by Saxena et al. (2017) more than 10 radio-loud AGN at $z > 6$ are expected to be detected in all the LOFAR Deep Fields combined with the current sensitivity. Currently, the most distant radio galaxy has a redshift of $z = 5.72$ (Saxena et al. 2018b). The SILVERRUSH catalogue is potentially an excellent source for finding radio galaxies at $z > 6$, even though only a small fraction of the volume is probed due to the detection range of the narrowband filters. The detection of such a high redshift radio galaxy would not only be a substantial step forward in the study of the formation and evolution of massive galaxies into the EoR, but would also be the most distant radio galaxy discovered to date.

This paper is structured as follows. In Sect. 2 we describe the details of the SILVERRUSH program, the LoTSS Deep Fields data used, and the LAE sample selection. In Sect. 3 we present the characteristics of the radio detected SILVERRUSH population. Subsequently, in Sect. 4 we analyse the optical/infrared properties and multi-wavelength stacks of the wider LAE population. In Sect. 5 we discuss the LOFAR detection rate and LAE selection. Finally, in Sect. 6 we summarise our findings.

In this work a flat Λ -CDM cosmology is assumed using $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$. Furthermore, all magnitudes presented are given in the AB system (Oke & Gunn 1983).

2. Data and sample selection

2.1. SILVERRUSH program

To identify LAEs, the SILVERRUSH survey uses 4 narrowband (NB) filters, which are mounted on the Hyper Suprime-Cam (HSC). These filters are NB387, NB816, NB921, and NB101, corresponding to central wavelengths (λ_{cen}) of 3858, 8169, 9204, and 10092 \AA^1 , which allow for identifying LAEs at redshifts of $z = 2.17 \pm 0.02$, 5.72 ± 0.05 , 6.57 ± 0.05 , and 7.30 ± 0.04 respectively, in addition to rest frame optical emission line galaxies at lower redshifts (see Aihara et al. 2018a). Using NB816 and NB921 imaging data of the HSC Subaru Strategic Program (HSC-SSP) survey, the SILVERRUSH program identified ~ 2000 LAE candidates at $z = 5.7$ and $z = 6.6$ in 5 fields with a total area of 14 and 21 deg^2 respectively (Shibuya et al. 2018b). These resulting LAE catalogues are publicly available² and will be referred to as NB816 and NB921 LAE catalogues in this work. The public data release does not include the results of LAE studies based on NB387 and NB101 observations. The detailed SILVERRUSH program strategy is described by Ouchi et al. (2018) and the first SILVERRUSH catalogues and properties are presented by Shibuya et al. (2018b). Other main scientific results following from SILVERRUSH, such as clustering properties and $\text{Ly}\alpha$ luminosity functions, are published in Shibuya et al. (2018a); Konno et al. (2018); Harikane et al. (2018); Inoue et al. (2018); Higuchi et al. (2019); Harikane et al. (2019); Kakuma et al. (2019). In the following sections we focus on the data and sample selection in the EN1 field. The data obtained by the HSC-SSP survey are outlined in Sect. 2.1.1 and the SILVERRUSH procedure for LAE selection from this survey is summarised in Sect. 2.1.2.

2.1.1. HSC-SSP survey

The HSC filter transmission curves of broadbands (BB) g , r , i , z , and y and narrowbands NB816 and NB921 are shown in Fig. 1, including the central wavelengths λ_{cen} and full width half maxima (FWHM) for the i and z band, and NB filters. The NB816 and NB921 LAE catalogues are created using HSC-SSP S16A data taken using these 5 BB filters and NB filters (Ouchi et al. 2018). Deep observations of EN1 have been conducted as part of the HSC-SSP survey, reaching 5σ limiting magnitudes of 25.3 for the NBs and ~ 24 –26 for the BBs (Shibuya et al. 2018b). The image reduction is performed using the HSC pipeline (Bosch et al. 2018). The source detection and photometric measurements are obtained using a ‘forced’ and ‘unforced’ method. In the unforced method, the coordinates, shape, and flux of each source are determined individually in each band, whereas in the forced method the coordinates and shape of the sources are fixed in a reference band and applied to all other bands to determine the flux. In this work, the photometry from the both forced and unforced method is used to ensure the largest possible sample. For further details on the HSC-SSP survey we refer to Aihara et al. (2018a) and Shibuya et al. (2018b).

¹ <http://svo2.cab.inta-csic.es/theory/fps/index.php?mode=browse&gname=Subaru&gname2=HSC>

² <http://cos.icrr.u-tokyo.ac.jp/rush.html>

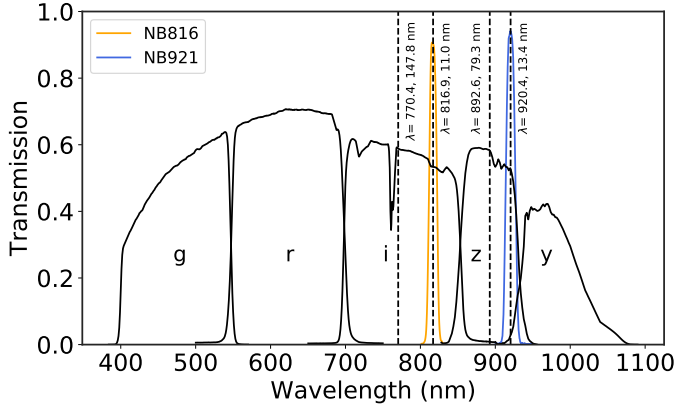


Fig. 1: Filter transmission curves of HSC BB (*grizy*) and NB (NB816 & NB921) filters. The central wavelength (λ_{cen}) and full width half maximum (in nm) are indicated for the *i*, *z*, NB816, and NB921 filters. The filter transmission curves and λ_{cen} values have been obtained from the SVO Filter Profile Service¹.

2.1.2. LAE selection

The LAEs in the SILVERRUSH program are selected using colour selection criteria (presented in Shibuya et al. 2018b and based on Ouchi et al. 2008, 2010), which ensure a non-detection in bands blue-ward of the Lyman break at a certain redshift and significant detected flux excess in the NBs. These criteria are given by:

$$\begin{aligned} i - \text{NB816} &\geq 1.2 \\ g &> g_{3\sigma} \end{aligned} \quad (1)$$

$$(r \leq r_{3\sigma} \text{ and } r - i \geq 1.0) \text{ or } (r > r_{3\sigma})$$

and

$$\begin{aligned} z - \text{NB921} &\geq 1.0 \\ g &> g_{3\sigma} \text{ and } r > r_{3\sigma} \end{aligned} \quad (2)$$

$$(z \leq z_{3\sigma} \text{ and } i - z \geq 1.3) \text{ or } (z > z_{3\sigma})$$

for $z = 5.7$ and $z = 6.6$ LAEs respectively, where $g_{3\sigma}$, $r_{3\sigma}$, and $z_{3\sigma}$ are the 3σ limiting magnitudes in the *g*, *r*, and *z* band. In the forced catalogue a stricter colour selection criterion of $z - \text{NB921} > 1.8$ is used and the limiting magnitudes are lowered to 2σ in the *g* and *r* band. These stricter colour criteria are also used in Konno et al. (2018) for studying the Ly α luminosity function. Besides the colour selection, several parameters and flags are used to denote imaging problems or artefacts, such as sources containing saturated pixels, and images with short exposure times (see Shibuya et al. 2018b; Aihara et al. 2018b). Further details on the SILVERRUSH LAE selection algorithm are presented in Shibuya et al. (2018b).

The final SILVERRUSH samples in EN1 contain 229 and 349 sources at $z = 5.7$ and 6.6 respectively in a 6 deg^2 area (Konno et al. 2018). The reliability of their LAE selection in all fields is checked using spectroscopic observations of a subsample of the selected LAEs. In total, redshifts of 96 of the HSC LAEs were spectroscopically confirmed and yielded a contamination rate of 0-30%, depending on the magnitude ranges (Shibuya et al. 2018a).

2.2. LOFAR radio observations of EN1

LoTSS aims to cover the whole northern hemisphere, reaching $\leq 100 \mu\text{Jy beam}^{-1}$ noise levels, of which a part of the data has been released in LoTSS-DR1 (Shimwell et al. 2019). LoTSS is being complemented by a series of LoTSS Deep Fields, which aim to ultimately cover a sky area of $\sim 50 \text{ deg}^2$ down to noise levels of $\sim 10 \mu\text{Jy beam}^{-1}$, probing the fainter and higher redshift radio population. The LoTSS Deep Fields first data release targets the fields Lockman hole, Boötes, and EN1, which are by design at declinations optimal for LOFAR observing sensitivity. These fields have already been observed extensively across optical and infrared (IR) wavelengths, enabling the determination of photometric redshifts and physical galaxy properties. New imaging and calibration algorithms have been developed by Tasse et al. (2020) to be able to construct thermal noise-limited images, and the LoTSS Deep Fields radio images reach $\sim 20 \mu\text{Jy beam}^{-1}$ RMS sensitivities in EN1 (Sabater et al. 2020). Details on the EN1 radio imaging and catalogue are presented in Sabater et al. (2020).

2.3. Deep optical/IR photometry in EN1

An extensive range of multi-wavelength observations from ultraviolet (UV) to far-infrared (FIR) are available in the EN1 field. An overview of the observations used in this work are given in Tab. 1. Here, we provide a summary of these observations and the generation of the multi-wavelength catalogue. We refer to Kondapally et al. (2020) for a detailed description of the catalogues used.

All LOFAR detected radio sources are cross-matched with the multi-wavelength catalogues, and photometric measurements extracted, by Kondapally et al. (2020) and the reader is referred to that paper for full details of the procedure. In summary, a multi-wavelength catalogue is created using forced, matched aperture photometry on pixel-matched images from all surveys. To achieve this, all individual images are resampled to a pixel scale of $0.2''$ and sky background is subtracted before adding them together using SWarp (Bertin et al. 2002). The flux is adjusted to a common scale using the zero-point magnitude, exposure time, and Vega-AB conversion factors, where needed. Optimal signal to noise (or χ^2) detection images are created to be able to detect the faintest sources. These are created by stacking multiple bands using SWarp for the optical to near-infrared (NIR) and Spitzer Infrared Array Camera (IRAC) observations separately, due to lower resolution of the Spitzer data. The weight assigned to each band in the χ^2 detection image varies according to the colour of the source. The sources are extracted from the χ^2 detection image using SExtractor (Bertin & Arnouts 1996) and the two source lists are combined to make a single catalogue. The fluxes in the different bands are obtained from all sources detected in either of the χ^2 detection images using aperture sizes with a diameter of $1''$ - $7''$ (in steps of $1''$) and $10''$. The varying point spread functions (PSF) in each filter is corrected for by aperture corrections, determined using the curve of growth estimated from moderately bright sources (see Kondapally et al. 2020). In this work we use the flux measurements from the $3''$ apertures for the optical-NIR filters and $4''$ apertures for the Spitzer-IRAC filters, which have both been aperture corrected. The $3''$ aperture for the optical-NIR filters is less affected by PSF variations than a $2''$ aperture and will therefore result in more robust colours (Kondapally et al. 2020). The final EN1 multi-wavelength catalogue contains over 2.1 million sources, of which 1.5 million sources in the overlapping

Table 1: Overview of the multi-wavelength observations in EN1 used for this work. The area covered by each survey and typical depths (3σ) per survey, except for the SPIRE, PACS, and MIPS, are adopted from [Kondapally et al. \(2020\)](#). See that paper for the detailed 3σ depth per filter. The depth for SPIRE, PACS, and MIPS are given by flux limits in mJy. For details on the optical Panoramic Survey Telescope and Rapid Response System (PanSTARRS) Medium Deep Survey, HSC-SSP survey, and Canada France Hawaii Telescope (CFHT) MegaCam SpARCS, see [Chambers et al. \(2016\)](#), [Aihara et al. \(2018b\)](#), and [Wilson et al. \(2009\)](#) respectively. A detailed description of UKIRT Infrared Deep Sky Survey - Deep Extragalactic Survey (UKIDSS-DXS) conducted with Wide Field Camera (WFCAM) on the United Kingdom Infra-Red Telescope (UKIRT) can be found in [Lawrence et al. \(2007\)](#). The Herschel Multi-tiered Extragalactic Survey (HerMES) is described in [Oliver et al. \(2012\)](#) and Spitzer Spitzer Wide-Area Infrared Extragalactic Survey (SWIRE) and Spitzer Extragalactic Representative Volume Survey (SERVS) in [Lonsdale et al. \(2003\)](#) and [Mauduit et al. \(2012\)](#) respectively.

Telescope	Instrument	Survey	Wavelength cov	Area (deg ²)	Depth
LOFAR		LoTSS Deep Fields	150 MHz	~25	~20 μ Jy beam ⁻¹
Herschel	SPIRE	HerMES	250, 350, 500 μ m	7.16	4 mJy
	PACS	HerMES	110, 170 μ m	7.13	12.5/17.5 mJy
Spitzer	MIPS	SWIRE	24 μ m	7.16	20 mJy
	IRAC	SWIRE	3.6, 4.5, 5.8, 8.0 μ m	9.32	~22.2 mag
		SERVS	3.6, 4.5 μ m	2.39	24.1 mag
UKIRT	WFCAM	UKIDSS-DXS	J and K	8.87	~23.0 mag
Haleakala Observatory	PanSTARRS	Medium Deep Survey	grizy	8.05	~24.7 mag
Subaru Telescope	HSC	SSP survey	grizy, NB816, NB921	7.70	~24.5 mag
CFHT	MegaCam	SpARCS	u	11.81	25.4 mag

PanSTARRS, UKIDSS-DXS, and Spitzer-SWIRE surveys are used for radio-optical cross-matching in an area of 6.7 deg². The cross-matching is carried out using an adaptation of the technique developed in LoTSS DR1 and presented by [Williams et al. \(2019\)](#). In short, optical/IR counterparts of radio sources are identified by either the statistical likelihood ratio method or by visual classification schemes, as determined by a decision tree described in [Williams et al. \(2019\)](#) and [Kondapally et al. \(2020\)](#). Sources with extended and/or complex radio emission are associated and classified using a combination of the LOFAR Galaxy Zoo ([Williams et al. 2019](#)) and an expert-user work-flow designed for deblending radio sources. For a detailed description of the procedure followed to create the catalogues we refer to [Kondapally et al. \(2020\)](#).

In addition, the Spectral and Photometric Imaging Receiver (SPIRE) and Photodetector Array Camera and Spectrometer (PACS) FIR fluxes are added to the catalogue obtained from the Herschel Extragalactic Legacy Project (HELP; [Oliver et al. 2020](#), in prep). The procedure for the determination of the FIR fluxes is described in [McCheyne et al. \(2020\)](#). In short, mid- and far-IR fluxes are derived from prior driven deblending of the Spitzer Multiband Imaging Photometer (MIPS) and Herschel PACS/SPIRE imaging using the `xid+` deblending code ([Hurley et al. 2017](#)).

Photometric redshifts for all optical sources in the [Kondapally et al. \(2020\)](#) catalogues were estimated using a hybrid template and machine-learning method. See [Duncan et al. \(2020a\)](#), hereafter D20) for full details of the method and characterisation of the redshift performance.

3. Characterising the radio detected SILVERRUSH population

The LOFAR detected SILVERRUSH population in EN1 is determined by cross-matching sources from the LoTSS catalogue and SILVERRUSH NB921 and NB816 LAE catalogues within 3" separation (half of the LOFAR 6" beam size). The location of optical counterparts of the radio sources are used for cross-matching. The resulting cross-matched sources are inspected by

eye, leading to 5 LOFAR detected SILVERRUSH sources out of the 8 potential matches, which are from now on referred to as the ‘LOFAR detected sample’. All of these 5 sources matched with SILVERRUSH sources in the NB921 LAE catalogue ($z = 6.6$), whereas no sources have been found to match with the NB816 LAE catalogue. The multi-wavelength cutouts of the LOFAR detected sample are shown in Fig. 2.

Tab. 2 summarises the photometric source properties including the photometric redshift (photo- z s) derived by D20. Only one source (ILTJ160658.74+550607.0) has a photo- $z > 6$, whereas estimates for other sources place them at $1.4 < z < 2.1$, suggesting these are likely to be low-redshift interlopers. It is furthermore noticeable that $z - \text{NB921} \geq 1.0$ within 1σ in the LoTSS photometry, but not as strong as the HSC values. This is likely due to the difference in underlying photometry used, and the large uncertainties in z magnitude, which could partly be due to the chosen 3" aperture size. To further investigate the nature of these 5 sources and their probable redshifts, we examine their spectral energy distributions (SEDs). In SED fitting we also include the FIR data, which is not included in the photo- z estimates.

3.1. SED fitting

We perform SED fitting on the LOFAR detected sample using the code Bayesian Analysis of Galaxies for Physical Inference and Parameter ESTimation³ (or BAGPIPES; [Carnall et al. 2018](#)). In summary, BAGPIPES is a Bayesian spectral fitting code developed for generating complex model galaxy spectra from photometric and spectroscopic observations using the MultiNest nested sampling algorithm ([Feroz et al. 2019](#)). The input observations can range from far-UV to microwave wavelengths. Model parameters used for fitting are star formation history (SFH), nebular component, and dust attenuation and emission components. An overview of the parameter values used in this work is given in Tab. A.1 and the input photometric data available are summarised in Tab. 1. The SFH is modelled using a double power

³ <https://bagpipes.readthedocs.io/en/latest/>

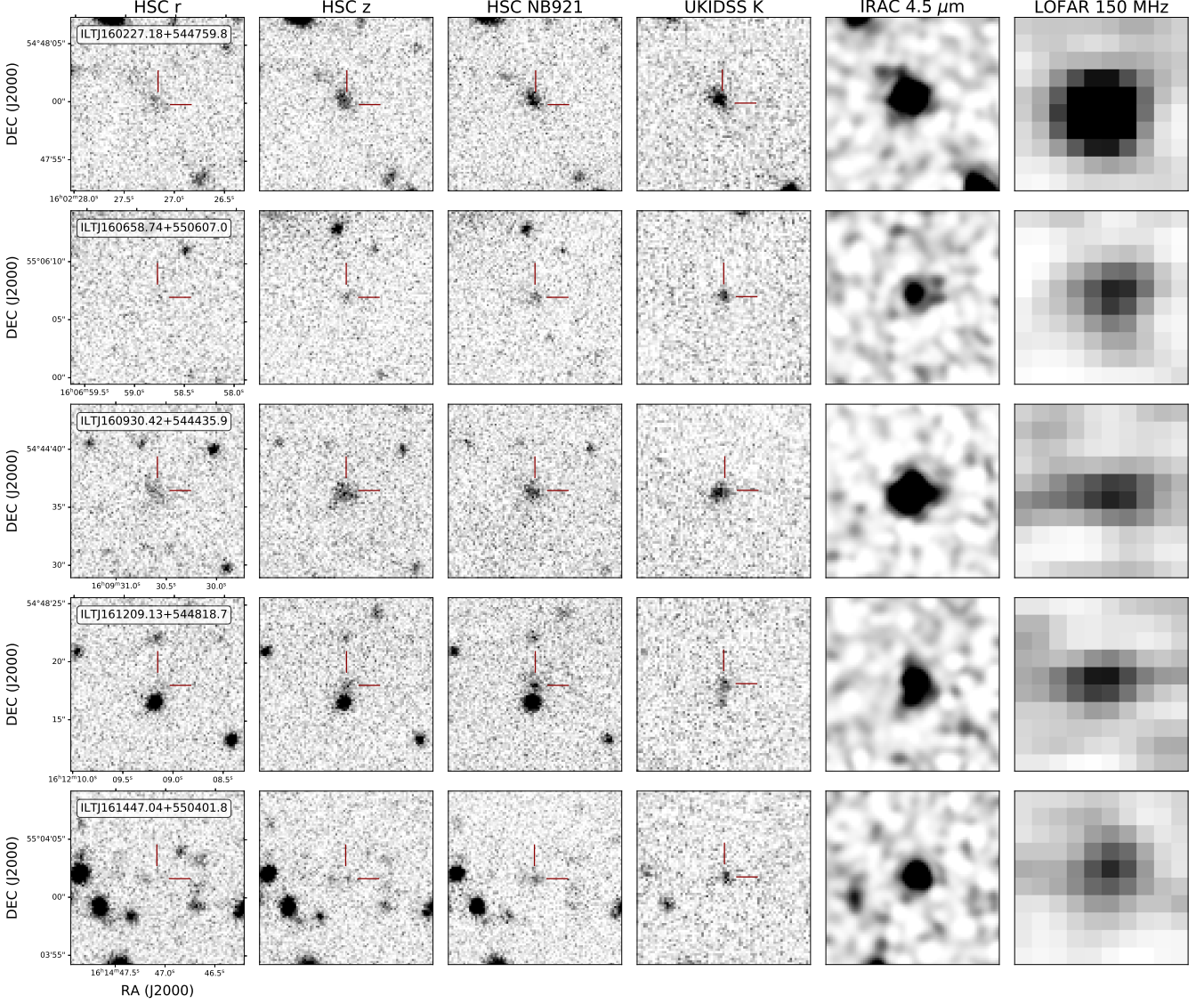


Fig. 2: Multi-wavelength cutouts of the LOFAR detected sample including from left to right: HSC r band, HSC z band, HSC NB921 band, UKIDSS K band, IRAC 4.5 μm , LOFAR 150 MHz. The sources show negligible or no emission below the Lyman break in the r band, however clear emission is seen in both UKIDSS k band and IRAC.

law with a parameter for the total stellar mass formed and metallicity, and the dust attenuation model of Calzetti et al. (2000) is used. For a further description of BAGPIPES and its performance we refer to Carnall et al. (2018).

Besides Ly α , other strong nebular emission lines that could account for a narrowband excess in the NB921 filter are H α , [OIII], H β , and [OII] nebular emission lines at $z \sim 0.40, 0.84, 0.89$, and 1.47 respectively. The photo- z accuracy is not high enough to be able to distinguish the H β and the [OIII] emission line, so throughout this work we refer to it as the [OIII] emission line at $z \sim 0.84$, since this line is more prevalent in emission line galaxies (e.g. Hayashi et al. 2018; Sobral et al. 2015a; Khostovan et al. 2015). The combination of implausibly high dust extinction required to produce such red SEDs at $z = 0.4$ and proportionally low co-moving volume probed means that contamination of this sample by H α emitters at $z = 0.4$ is thought

to be negligible. Therefore, specific models for [OIII], [OII], and Ly α emission lines are fitted to the LAE candidates.

It must be noted that BAGPIPES does not include AGN templates, which could lead to an underestimation of the mid-IR flux. An AGN contribution can be therefore identified by the inability of BAGPIPES to fit the mid-IR flux. The possible effects of the lack of AGN templates are therefore discussed together with the results.

To investigate the probability of a source residing at the above mentioned redshifts the SED fits are performed with fixed $z = 0.84, 1.47, 6.57$ and also at 80 redshifts in the range $z = [0.1, 8]$, in steps of 0.1. For each redshift, the SED fit resulting in the minimum reduced chi-squared value (χ^2_ν) is determined, to be used as a comparison for the different models. Also non-detections, with $\text{SNR} < 3.0$, are used by BAGPIPES and in the χ^2_ν calculation. However, negative fluxes are not taken into account. The SED fitting results of the high- z LAE candidate

Table 2: Multi-wavelength photometric properties and photo-zs of LOFAR selected sample. The $4.5 \mu\text{m}$ magnitude is taken from the SWIRE survey (Lonsdale et al. 2003) and the i and z magnitude from HSC-SSP (Aihara et al. 2018b), both from the LoTSS Deep Fields catalogue (Kondapally et al. 2020). The photometric redshifts $z_{1,\text{median}}$ and $z_{2,\text{median}}$ are obtained from LoTSS Deep Fields catalogue (see D20), where $z_{1,\text{median}}$ and $z_{2,\text{median}}$ are the medians of the primary and secondary redshift peak, respectively. More detailed SED fitting shows that these sources are all likely [OIII] emitters at $z=1.47$ (see Sect. 3.1).

Source name	$S_{150\text{MHz}}$ (mJy)	$4.5 \mu\text{m}$ mag	K mag	NB921 mag	i mag	z mag	photo- z	
							z_1	z_2
ILTJ160227.18+544759.8	0.35 ± 0.04	19.95 ± 0.07	21.10 ± 0.11	23.18 ± 0.15	24.64 ± 0.43	23.76 ± 0.32	$1.7^{+0.8}_{-0.6}$	$2.7^{+0.1}_{-0.1}$
ILTJ160658.74+550607.0	0.10 ± 0.04	20.99 ± 0.17	22.39 ± 0.34	24.48 ± 0.50	25.60 ± 1.04	25.33 ± 1.37	$6.1^{+0.5}_{-0.8}$	$4.4^{+0.6}_{-0.6}$
ILTJ160930.42+544435.9	0.11 ± 0.05	20.13 ± 0.08	21.17 ± 0.12	23.19 ± 0.15	24.59 ± 0.41	23.90 ± 0.37	$1.6^{+0.5}_{-0.4}$	-
ILTJ161209.13+544818.7	0.18 ± 0.05	20.57 ± 0.12	21.65 ± 0.18	22.26 ± 0.06	23.49 ± 0.15	23.16 ± 0.19	$1.4^{+0.1}_{-0.2}$	$0.85^{+0.06}_{-0.03}$
ILTJ161447.04+550401.8	0.09 ± 0.03	20.73 ± 0.13	22.16 ± 0.29	24.49 ± 0.50	25.76 ± 1.19	25.71 ± 1.94	$2.1^{+1.4}_{-1.0}$	$4.4^{+0.6}_{-0.6}$

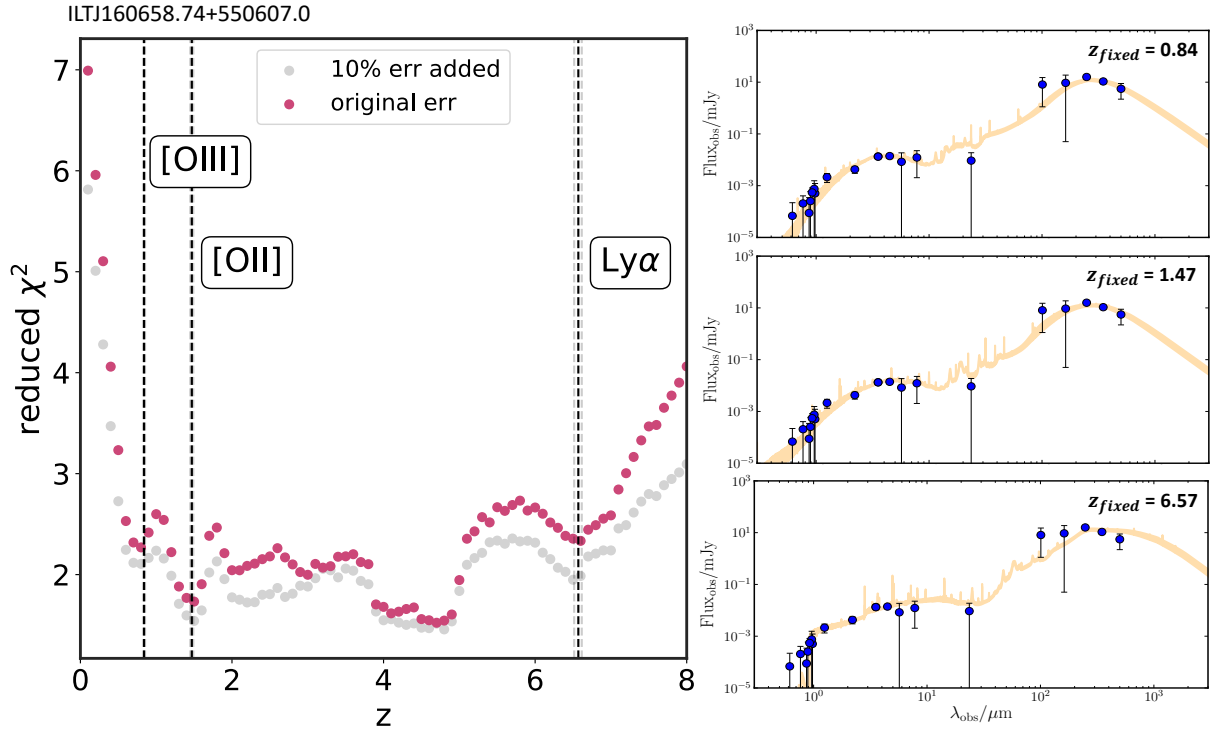


Fig. 3: SED fitting results for ILTJ160658.74+550607.0, obtained using BAGPIPES. Left: Reduced χ^2 for SED fits in the range $0 < z < 8$, with locations of $\text{Ly}\alpha$, [OII], and [OIII]+H β emission line indicated by black dashed lines. The grey dashed lines indicate the redshift range probed by the NB filter for the strong emission lines [OIII], [OII], as well as $\text{Ly}\alpha$. Right: observed flux of ILTJ160658.74+550607.0 (blue) in mJy with the original errors and posterior median model (yellow) for fits with redshifts fixed at $z = 0.84, 1.47$, and 6.57 . The source is best fitted by a template with $z=1.47$ when considering the possible emission lines. No other emission line is strong enough for the source to be at $z \sim 4$.

ILTJ160658.74+550607.0 are shown in Fig. 3. The χ^2_ν value lies within a small range and is lowest at $z \sim 4.9$ when fitting with the original errors. However, in panchromatic SED fitting, there are many model uncertainties and flux errors arising from flux calibration issues or correlated noise are usually underestimated (see e.g. Marquez et al. 2014). We therefore re-fit the LOFAR detected sample incorporating an additional 10 % flux uncertainty in quadrature to assess what effect this may have on the preferred redshift solution (see Fig. 3). The overall shape of the curve remains largely the same with the additional 10% flux uncertainty, including the decrease around $z \sim 4 - 5$, but it does improve the fit at $z = 1.47$, making it comparable within the errors to the $z \sim 4 - 5$ fit. However, around this redshift of $z \sim 4 - 5$ there is no emission line with a typical equivalent width high

enough to be detected in the HSC NB921 filter. Also, because of the large number of data points, the relative contribution of the narrowband measurement to the overall χ^2 value is limited. The decrease in χ^2 can therefore be attributed to the broadband SED shape. In the right panel of Fig. 3 the best-fitting posterior median models at $z = 0.84, 1.47$, and 6.57 for [OIII], [OII], and $\text{Ly}\alpha$ respectively are shown in yellow. Here, the SED fit at $z = 1.47$ is the preferred solution. The SED fitting results of the other 4 sources in the LOFAR detected sample are shown in Fig. B.1. The χ^2 values are minimised around $z \sim 1.5$, providing strong evidence that these are [OIII] emitters at $z = 1.47$.

To compare the preference for each model, we use the Bayesian Information Criterion (BIC), which in its general form

is given by

$$\text{BIC} = -2 \ln L + k \ln N, \quad (3)$$

where L is the maximised likelihood, k the number of parameters, and N the number of observations (see Kass & Raftery 1995). In the case of Gaussian distributed model errors this becomes

$$\text{BIC} = \chi^2 + k \ln N. \quad (4)$$

The model with the highest probability minimises this value and when comparing models a $\Delta\text{BIC} > 2$ gives positive evidence for one of the models being preferred over the other, whereas $-2 < \Delta\text{BIC} < 2$ indicates that there is no evidence for either model being preferred. A $\Delta\text{BIC} > 6$ gives a strong positive evidence (as defined by Kass & Raftery 1995). When comparing our SED fits, the number of observations and parameters remains constant, therefore BIC becomes χ^2 and we can use the $\Delta\chi^2$ values of the fits at $z = 0.84, 1.47$, and 6.57 , to determine the preference for each model. The χ^2_ν and $\Delta\chi^2$ values of the Ly α , [OII], and [OIII] emission lines fits to the LOFAR selected sample are summarised in Tab. 3. Here, the original errors are used for comparison of the models, since the 10% added flux error causes over-fitting of the data resulting in χ^2_ν values below 1. For all 5 sources, $\Delta\chi^2(\text{Ly}\alpha - [\text{OII}])$ and/or $\Delta\chi^2(\text{Ly}\alpha - [\text{OIII}]) > 2$ meaning the models for [OII] or [OIII] provides a better fit to the data and Ly α is likely not to be the detected emission line. In addition, $\Delta\chi^2([\text{OIII}] - [\text{OII}]) > 2$ for all sources, suggesting the detected emission line is [OII], implying the sources are situated at $z = 1.47$. This conclusion is also supported by comparing the co-moving volume being probed at $z=0.84$ and $z=1.47$, which is ~ 2 times higher for $z = 1.47$ when assuming a flat Λ -CDM cosmology. The evidence for Ly α not being the source of emission is less strong in the case of ILTJ160658.74+550607.0 (shown in Fig.3), however the FIR SED shape, consisting of SPIRE and PACS observations, is best reproduced by the $z = 1.47$ model and therefore seems to be the actual redshift. Follow-up observations are needed to confirm the redshift of this source. If this source is a radio galaxy at $z = 6.6$, it would be the most distant radio galaxy known to date.

We present the derived stellar mass and star formation rate for each fixed-redshift SED fit in Tab. 4. When redshift is fixed at $z = 6.57$, the derived stellar masses are $10^{11.7} < M_* < 10^{11.9} M_\odot$ and SFRs $\sim 650\text{--}4400 M_\odot \text{ yr}^{-1}$. Even given the large uncertainties on the high mass end of the stellar mass function at these redshifts, the likelihood of masses being this high is small (see e.g. Duncan et al. 2014; Grazian et al. 2015; Song et al. 2016). The median Ly α SFR obtained by e.g. Calhau et al. (2020) is $9.8^{+9.7}_{-5.2} M_\odot \text{ yr}^{-1}$ for their sample at $z > 3.5$, when excluding AGNs. The SFRs derived from their FIR detected LAEs (of which 76% are X-ray or radio AGN) using their IR fluxes yield higher values of $200^{+430}_{-110} M_\odot \text{ yr}^{-1}$, with a few sources $> 600 M_\odot \text{ yr}^{-1}$.

As noted before, BAGPIPES does not include AGN templates. If there was a strong AGN contribution, there should be a warm dust contribution at mid-IR wavelengths. However, none of these 5 sources show excess in the MIPS $24\mu\text{m}$ band. In addition, all these sources are classified as star-forming from both their optical and radio emission via the hybrid SED fitting method of Best et al. (2020). Furthermore, in the photo- z fitting procedure of D20 AGN templates are used, and these results are consistent with the redshifts obtained from SED fitting. Unlike in D20, in the BAGPIPES modelling FIR fluxes have been included in the fit, which for the source ILTJ160658.74+550607.0 helps break the degeneracy between the $z = 1.5$ and 6.6 solution. The remaining

small differences in resulting redshifts between this work and the photo- z s of D20 can be attributed to the use of a different set of templates (see Duncan et al. 2018) and the machine learning contribution in D20.

3.2. Emission line and radio properties

From the photometric data, physical properties of the LOFAR selected sample, such as radio luminosity, line luminosity, and equivalent width (EW) are determined (see Tab. 5). As each of these quantities depend on the redshift of the source, the values are given for fixed $z = 0.84, 1.47$, and 6.57 . A spectral index α of -0.7 ± 0.7 (defined by $S_\nu \propto \nu^\alpha$ with S_ν the flux density in Jy) is assumed for determination of the radio luminosity, where the error is based on the spectral index distribution derived by Saxena et al. (2018a). No spectral index could be obtained for the sources in the LOFAR selected sample, since no sufficiently deep radio data (e.g. from the VLA or GMRT) was publicly available. For calculation of the EW and line luminosity, we use the equations presented in Sobral et al. (2012). Subsequently, the line luminosity and radio luminosity are calculated using the line flux and radio flux respectively. The rest-frame equivalent width (EW_0) is determined by dividing the observed EW by $1 + z$ to correct for the Hubble expansion. The values obtained for each assumed redshift are shown in Tab. 5.

A large range of Ly α rest EW values are quoted in literature. Ly α EW values of e.g. $\sim 0\text{--}400 \text{ \AA}$ (Ono et al. 2010) and $10\text{--}1000 \text{ \AA}$ (Leclercq et al. 2017) have been found. Furthermore, the Ly α luminosity functions (LF) presented in Calhau et al. (2020) and Konno et al. (2018) show Ly α luminosities ranging from $\sim 10^{42.5} \text{--} 10^{44.0} \text{ erg s}^{-1}$. A global study of LAEs at $2.5 < z < 5.8$ by Sobral et al. (2018b) in the COSMOS field resulted in a characteristic luminosity $L_{\text{Ly}\alpha} \sim 10^{43.8 \pm 0.1} \text{ erg s}^{-1}$ of the Schechter function. The values for the EW and $L_{\text{Ly}\alpha}$ found in this work are consistent with these literature results.

The EW_0 derived when we assume that the detected lines are [OII] and [OIII] are comparable to the EW distribution found in Khostovan et al. (2016), where the majority of the [OIII]+H β line emitters at $z = 0.84$ have $20 < \text{EW}_0 < 500$ and the majority of [OII] line emitters at $z = 1.47$ have $20 < \text{EW}_0 < 300$. The derived [OII] and [OIII] luminosities are within the ranges of $41.4 < \log_{10}(L_{[\text{OII}]}) < 42.3$ and $40.5 < \log_{10}(L_{[\text{OIII}]}) < 42.1$. However, it must be noted that the sources in this study could be different from the general population of [OII] and [OIII] emitters, since this sub sample of sources satisfy the LAE selection criteria. These criteria (see eq. 1 and 2) select for sources with high magnitude differences between the z band and NBs, and non-detections in filters below the supposed Lyman break, suggesting these are potentially sources with strong Balmer breaks.

The obtained radio luminosities and SFRs from BAGPIPES can be compared to the $L_{150\text{MHz}}$ -SFR relation derived by Gürkan et al. (2018), assuming the relation does not evolve with redshift. This assumption is supported by the study of Duncan et al. (2020b), where no redshift evolution is found out to $z \sim 2.6$. If the sources are located at $z = 1.5$, all 5 sources are slightly offset to higher radio luminosities than would be expected from the $L_{150\text{MHz}}$ -SFR relation, with a difference of $0.04\text{--}0.69$ dex from the 1σ upper limit, which indicates they could be low luminosity AGN. However, these deviations are non-significant when taking into account the 0.3 dex error on the radio luminosities and ~ 0.3 dex scatter of observations in the $L_{150\text{MHz}}$ -SFR relation. Moreover, the SED fitting did not suggest an AGN contribution. The Gürkan et al. (2018) relation is consistent with the new derived

Table 3: Reduced χ^2 and $\Delta\chi^2$ values from SED fitting (using original error) to redshift $z = 6.57 \pm 0.05$, 1.47 ± 0.02 , 0.84 ± 0.02 corresponding to emission lines Ly α , [OII], and [OIII] respectively taking into account the FWHM (133.45 Å) of the NB921 HSC filter. The $\Delta\chi^2$ values suggest that the 5 sources are [OII] emitters at $z = 1.47$.

Source name	Reduced χ^2			$\Delta\chi^2$		
	Ly α	[OII]	[OIII]	Ly α – [OII]	Ly α – [OIII]	[OIII] – [OII]
ILTJ160227.18+544759.8	5.0 ^{+3.1} _{-0.6}	0.9 ^{+0.0} _{-0.1}	2.0 ^{+0.2} _{-0.0}	53-95	33-77	18-21
ILTJ160658.74+550607.0	2.3 ^{+0.2} _{-0.0}	1.7 ^{+0.1} _{-0.1}	2.3 ^{+0.1} _{-0.0}	5.7-8.6	0.1-2.0	5.8-7.5
ILTJ160930.42+544435.9	4.6 ^{+1.69} _{-0.2}	0.9 ^{+0.1} _{-0.1}	2.1 ^{+0.1} _{-0.1}	60-93	38-73	19-23
ILTJ161209.13+544818.7	25.6 ^{+3.3} _{-2.7}	0.6 ^{+0.5} _{-0.0}	5.5 ^{+0.2} _{-0.3}	371-476	293-401	71-83
ILTJ161447.04+550401.8	1.1 ^{+0.2} _{-0.1}	0.4 ^{+0.0} _{-0.0}	0.79 ^{+0.0} _{-0.05}	8.4-11.4	3.5-7.1	4.3-4.9

Table 4: Median (50th percentile) stellar mass and SFR obtained from SED fitting of the LOFAR selected sample, with fixed redshifts at $z = 0.84$, 1.47 , and 6.57 . The errors are given by the 84 and 16 percentiles.

Source name	$z = 0.84$		$z = 1.47$		$z = 6.57$	
	$\log M_*$ (M_\odot)	SFR ($M_\odot \text{ yr}^{-1}$)	$\log M_*$ (M_\odot)	SFR ($M_\odot \text{ yr}^{-1}$)	$\log M_*$ (M_\odot)	SFR ($M_\odot \text{ yr}^{-1}$)
ILTJ160227.18+544759.8	10.61 ^{+0.09} _{-0.11}	18 ⁺⁵ ₋₂	11.12 ^{+0.08} _{-0.27}	82 ⁺²⁶ ₋₁₅	11.78 ^{+0.18} _{-0.14}	4.4 ^{+0.8} _{-0.6} $\times 10^3$
ILTJ160658.74+550607.0	10.38 ^{+0.12} _{-0.18}	10 ⁺² ₋₁	10.75 ^{+0.18} _{-0.17}	54 ⁺¹³ ₋₁₀	11.69 ^{+0.20} _{-0.23}	2.2 ^{+0.3} _{-0.3} $\times 10^3$
ILTJ160930.42+544435.9	10.65 ^{+0.04} _{-0.09}	15 ⁺¹ ₋₂	11.13 ^{+0.03} _{-0.06}	61 ⁺⁸ ₋₈	11.87 ^{+0.24} _{-0.12}	2.6 ^{+0.4} _{-0.3} $\times 10^3$
ILTJ161209.13+544818.7	10.25 ^{+0.04} _{-0.06}	6 ⁺¹ ₋₁	10.70 ^{+0.04} _{-0.08}	24 ⁺⁴ ₋₃	11.68 ^{+0.06} _{-0.07}	8.8 ^{+0.2} _{-0.2} $\times 10^2$
ILTJ161447.04+550401.8	10.38 ^{+0.14} _{-0.18}	10 ⁺⁶ ₋₂	10.79 ^{+0.10} _{-0.12}	31 ⁺¹¹ ₋₇	11.91 ^{+0.13} _{-0.13}	6.5 ^{+0.7} _{-0.6} $\times 10^2$

$L_{150\text{MHz}}$ -SFR relation by Smith et al. (2020) from the LOFAR Deep Fields data.

4. Radio and optical/infrared properties of the wider LAE population

The evidence available suggests that the LAE candidates in the LOFAR selected sample are not located at $z \sim 6.6$. Therefore, we cannot give an upper limit on the AGN fraction of LAE at $z > 6$. These results however highlight a different possible problem of impurity of LAE samples and the low reliability of measurements on the LAE properties and the claimed AGN fractions at these redshifts. To investigate the scale of this problem, we further analyse the wider $z \sim 5.7$ and $z \sim 6.6$ LAE samples using our full panchromatic photometry.

The compiled optical/IR photometric catalogue for EN1 also allows for the study of the multi-wavelength properties of SILVERRUSH sources that do not have radio detections. The sources from the LoTSS multi-wavelength catalogue and SILVERRUSH catalogues are cross-matched within 1" and inspected by eye. Of the 349 and 229 LAEs in the SILVERRUSH NB921 and NB816 LAE catalogues in EN1, 58 and 53 sources, respectively, are detected in the LoTSS multi-wavelength catalogue. This sample will from now on be referred to as the ‘optically-selected sample’. The LoTSS catalogue, selected from the χ^2 detection image, may not be as optimised for Ly α candidates as the SILVERRUSH extraction, and SILVERRUSH used the deeper internal intermediate data release HSC-SSP 16A, so faint sources ($i \text{ mag} > 26$) are often missed in the LoTSS catalogue. We performed forced photometry on all SILVERRUSH sources to compare the two samples. The K and NB magnitudes of the optically-selected and non-detected SILVERRUSH sample are shown in Fig. 4. The non-detected SILVERRUSH sources with a flux measurement at $> 1\sigma$ have fainter K band magnitudes ($K = 23.0/23.2$ for NB816/NB921) compared to optically-

selected sources ($K = 22.1/22.3$ for NB816/NB921). The median NB magnitudes are only slightly fainter for the non-detected SILVERRUSH sources ($NB = 24.2/24.6$ for NB816/NB921) compared to optically-selected sources ($NB = 23.9/24.4$ for NB816/NB921). This indicates that the optically-selected sample contains sources which are brighter and redder than the non-detected SILVERRUSH sample. Note that these magnitudes differences are estimates, because of the faintness of the sources and the choice of a 1σ detection limit. The bias of our sample is further discussed in Sect. 5.2.

Properties of the optically-selected sample obtained from photometry and SED model fitting are discussed in Sect. 4.1. Furthermore, we stack the radio and infrared and optical images of the full sample to determine the limiting magnitudes, which are presented in Sect. 4.2.

4.1. SED fitting optically-selected sample

The stacked full photometric redshift probabilities from D20 for the optically-selected sample are shown in the left panel of Fig. 5. This is a sum of the photo-zs of the individual sources. Of the two probability distributions, 44% of the posterior is at $z > 5$ for the NB816 sources, and less than 10% of the posterior is at $z > 6$ for the NB921 sources. This already suggests a high contamination fraction in both of these samples.

These photo-zs could potentially be biased towards lower redshifts due to the applied magnitude prior, which favors low- z solutions ($z < 2$) even for faint sources, while not disallowing high- z solutions (Duncan et al. 2018). We therefore extend this with more detailed SED fitting by repeating the procedure in Sect. 3.1 for the full optically-selected sample of LAEs by fixing the redshifts at $z = 0.63, 1.19, 5.72$ for the NB816 catalogue and $z = 0.84, 1.47, 6.57$ for the NB921 catalogue. Of the 58 candidate LAEs at $z = 5.7$, 9 have FIR detections. For the $z = 6.6$ candidates, 6 out of 53 have FIR detections. The other

Table 5: Derived total radio luminosity $L_{150\text{MHz}}$, line luminosity L , and rest-frame EW_0 of the LOFAR selected sample, for SED fits with redshift fixed at $z = 0.84, 1.47$, and 6.57 , using the LOFAR 150 MHz total flux measurements and HSC flux values in the LoTSS catalogue.

Source name	$z = 0.84$			$z = 1.47$			$z = 6.57$		
	$\log_{10} L_{150\text{MHz}}$ (W Hz $^{-1}$)	$\log_{10}(L_{[\text{OIII}]})$ (erg s $^{-1}$)	$\text{EW}_{0,[\text{OIII}]}$ (Å)	$\log_{10} L_{150\text{MHz}}$ (W Hz $^{-1}$)	$\log_{10}(L_{[\text{OIII}]})$ (erg s $^{-1}$)	$\text{EW}_{0,[\text{OIII}]}$ (Å)	$\log_{10} L_{150\text{MHz}}$ (W Hz $^{-1}$)	$\log_{10}(L_{\text{Ly}\alpha})$ (erg s $^{-1}$)	$\text{EW}_{0,\text{Ly}\alpha}$ (Å)
ILTJ160227.18+544759.8	24.0 \pm 0.2	41.1 \pm 0.2	71 \pm 21	24.6 \pm 0.3	41.7 \pm 0.2	53 \pm 16	26.0 \pm 0.6	43.3 \pm 0.2	17 \pm 5
ILTJ160658.74+550607.0	23.5 \pm 0.3	40.7 \pm 0.6	132 \pm 151	24.0 \pm 0.3	41.3 \pm 0.6	99 \pm 113	25.4 \pm 0.6	42.9 \pm 0.6	32 \pm 36
ILTJ160930.42+544435.9	23.5 \pm 0.3	41.2 \pm 0.2	96 \pm 27	24.0 \pm 0.3	41.8 \pm 0.2	72 \pm 20	25.5 \pm 0.6	43.4 \pm 0.2	23 \pm 7
ILTJ161209.13+544818.7	23.7 \pm 0.2	41.6 \pm 0.1	152 \pm 26	24.3 \pm 0.3	42.2 \pm 0.1	114 \pm 19	25.7 \pm 0.6	43.8 \pm 0.1	37 \pm 6
ILTJ161447.04+550401.8	23.4 \pm 0.2	40.8 \pm 0.8	304 \pm 899	24.0 \pm 0.3	41.4 \pm 0.8	226 \pm 670	25.4 \pm 0.6	43.0 \pm 0.8	74 \pm 219

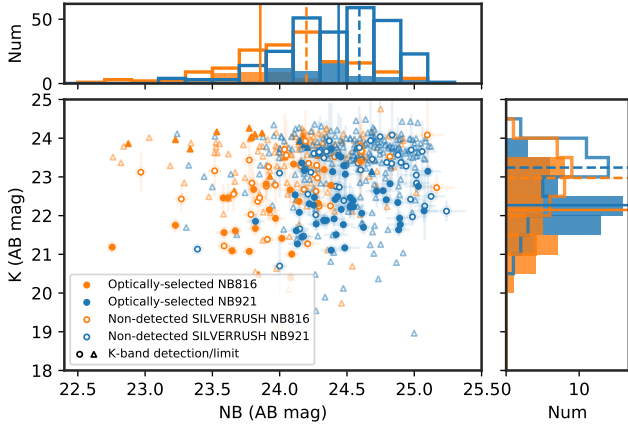


Fig. 4: K and NB magnitudes (NB816 or NB921) of optically-selected sample and non-detected SILVERRUSH sample. Limiting magnitudes (1σ) are given for sources below 1σ detection (filled and unfilled triangles for sources in the optically-selected and non-detected SILVERRUSH sample respectively). The histograms include only sources with a $>1\sigma$ detection. The median K and NB magnitudes are indicated by the normal and dashed lines for the optically-selected and non-detected sample respectively. In the optically selected sample 23% of sources are not detected in K above 1σ , while for the non-detected SILVERRUSH sample this fraction is 82%. The optically-selected sample is brighter and redder than the original SILVERRUSH sample.

sources have only UV-NIR data. We fit all available data for each source. For each source, we identify the strongest emission line with the best fitting SED (lowest χ^2_{ν}), and assign the sources the corresponding redshift. We present numbers of sources at each of the redshifts in Fig. 5 (right-hand panel). For 5 sources in NB921 and 1 source in the NB816 LAE catalogue, the SED fits are poorly constrained due to detections less than in 9 photometric bands, hence these results are excluded from further analysis. The best SED fits (right-hand panel in Fig. 5) suggest contamination rates of 90 and 93% for the NB816 and NB921 catalogue matches respectively. To quantify this further, we repeat the model comparison method for this larger non-radio-detected sample, again with fixed redshifts. The resulting $\Delta\chi^2$ values for $\Delta\chi^2(\text{Ly}\alpha - [\text{OIII}])$, $\Delta\chi^2(\text{Ly}\alpha - [\text{OII}])$, and $\Delta\chi^2([\text{OIII}] - [\text{OII}])$ are presented in Fig. 6, where the dotted line indicates a $\Delta\chi^2$ value of 2 and -2 . Sources likely to be Ly α emitters at $z = 6.57/5.72$ satisfy $\Delta\chi^2 < -2$ for either $\Delta\chi^2(\text{Ly}\alpha - [\text{OIII}])$ or $\Delta\chi^2(\text{Ly}\alpha - [\text{OII}])$, under the condition that the other $\Delta\chi^2 \not\geq 2$. This is the case for 4 sources in the NB816 and 5 sources in the NB921 catalogue. On the other hand, 43 and 47 of the sources have $\Delta\chi^2 > 2$ for either of the two $\Delta\chi^2$ values, for both NB816 and NB921, so

these are more likely to be $z = 0.63/0.84$ or $z = 1.19/1.47$, respectively, contaminants rather than $z = 5.72/6.57$ LAEs. The model comparison for sources with $\Delta\chi^2$ between -2 and 2 is inconclusive and can be regarded as an estimate of the uncertainty. The contamination rate determined from SED fitting using the three emission lines [OIII], [OII], and Ly α is therefore 81-92% (43-49 out of the 53 sources) and 81-91% (47-53 out of the 58 sources) for NB816 and NB921 respectively. Similarly, comparing the χ^2 of [OIII] and [OII], [OIII] is likely ($\Delta\chi^2 > 2$) to be the emission line for 63 \pm 31% and 55 \pm 28% of the possible contaminating sources for NB816 and NB921, respectively, whereas [OII] is likely to be the corresponding emission line for 6 \pm 31% and 17 \pm 28% of the sources for NB816 and NB921, respectively (with the remainder inconclusive). Comparing these results to the primary and secondary photo-zs of D20, we obtain that within the error, 21/53 (40%) of the photo-zs in the NB816 sample and 41/58 (71%) of the photo-zs in the NB921 sample are consistent. The remaining discrepancy between the two photo-z estimates can be attributed to the faint magnitudes of the sources and the high redshifts, which significantly increases the photo-z uncertainties and results in broad photo-z posterior (see D20). Comparing our contamination rates to the stacked photo-z probability distribution in the left panel of Fig. 5, shows that the photo-z estimates are consistent for the NB921 sample, but contamination is estimated to be lower for the NB816 sample. The photometric redshifts of D20 and SED fitting results together therefore suggest the contamination of our subset is 56-92% and 81-91% for the NB816 and NB921 sample, respectively.

Physical properties of the samples, derived from SED fitting (SFR and stellar mass) and from the photometric data (line luminosity and EW) are shown in Fig. 7 for each of the three proposed redshifts. Assuming these sources are LAE, the median Ly α luminosity is $10^{43.1\pm0.2}$ and $10^{43.0\pm0.2}$ erg s $^{-1}$ for NB816 and NB921, in line with other work of e.g. Calhau et al. (2020) and Konno et al. (2018). Fixing redshifts at $z = 1.47/1.19$ and $z = 0.84/0.63$ gives [OIII] and [OII] line luminosities that are also in good agreement with Khostovan et al. (2016) and Hayashi et al. (2018), as described in Sect. 3.2.

The EW_0 and SFR values derived for all three redshifts lie within the range of reported values for other samples in the literature. However, these physical parameters vary substantially both in this work and the literature, which is likely due to a wide range of galaxy types and survey selection functions (see e.g. Khostovan et al. 2019 and Calhau et al. 2020). If we assume the candidates to be low- z interlopers, we are again studying a subset of sources of the [OIII] and [OII] emitter population, which is likely biased, and therefore these sources do not necessarily have the same physical properties. We therefore refrain from making any further comparison of the obtained EW_0 and SFRs to support the evidence for the low- z nature of a high fraction of these sources. Remarkable, however, are the stellar masses for fixed

$z = 6.57$, which are extremely high ($M_* > 10^{11} M_\odot$) for especially the sources in the NB921 catalogue. The majority of these sources are therefore unlikely to be LAEs.

The derived source properties are also subject to systematic effects from the SED fitting models used and the possible biases introduced by the sample selection. BAGPIPES is constrained by the set input parameters, prior distributions, and model dependencies. Despite the limitations, the models can often successfully recreate the observed SED. However, spectroscopic observations are necessary to confirm the redshifts and derived physical properties.

4.2. Radio and IR/optical stacking

To investigate the radio properties of the LAE candidates not detected by LOFAR, first we stack the LOFAR images to set observational constraints on the average radio properties of the samples. SILVERRUSH sources outside the area covered by the multi-wavelength LoTSS data (see Sect. 2.3) are removed from the catalogue (22 and 3 sources for NB921 and NB816 catalogues, respectively). The 5 sources in the LOFAR detected sample in the NB921 catalogue are also removed from the stacking data set (see Sect. 3). For both the NB816/NB921, first the full sample is stacked (containing 226/322 sources) and then the optically-selected samples (containing 53/49 sources) are stacked separately, and these results are shown in the top and middle panels of Fig. 8. We assume the sources are unresolved in the 6" resolution radio images, and can therefore obtain the median flux from the peak pixels. The average RMS noise level for individual radio images is $\sim 25 \mu\text{Jy beam}^{-1}$. The stacked radio flux density, RMS, and resulting SNR obtained in each sample are summarised in Tab. 6. None of the stacked radio flux densities are significant (for all $\text{SNR} < 2$). Therefore, we can place a 3σ upper limit on the radio flux density of $\sim 4.8 \mu\text{Jy}$ for the sources in the full sample and $\sim 12.0 \mu\text{Jy}$ for sources in the optically-selected sample.

Furthermore, we stack the 4 and 5 optically-selected sources, in NB816 and NB921 respectively, that seem likely LAEs from SED fitting together with the 173 and 274 sources in the non-matched SILVERRUSH samples, under the assumption that they are LAEs. The non-matched sources are fainter and bluer and therefore less likely to be low-redshift contaminants (see Fig. 4 and Sect. 5.2). The stacks again yield no detection with a limiting radio flux density of $5.7 \mu\text{Jy}$ and $4.8 \mu\text{Jy}$ (3σ) for NB816 and NB921, respectively (see bottom panels Fig. 8). When assuming their redshifts to be $z = 5.7$ and 6.6 , the radio flux density can be converted to 2σ radio luminosity upper limits of 8.2×10^{23} and $8.7 \times 10^{23} \text{ W Hz}^{-1}$ respectively. Using the low- z $L_{150\text{MHz}}$ -SFR relation from Gürkan et al. (2018), an estimated 2σ upper limit on the SFR can be placed on these LAEs. Our derived upper limits are ~ 53 and $56 M_\odot \text{ yr}^{-1}$ on the $z = 5.7$ and $z = 6.6$ LAEs, respectively. The currently known relations between $L_{150\text{MHz}}$ and SFR are derived from low- z galaxies, therefore this is our current best SFR estimate. However, the work of Smith et al. (2020) suggests there is no strong redshift evolution. Similar SFRs are obtained with the $L_{150\text{MHz}}$ -SFR relations from Wang et al. (2019), with differences of a few $M_\odot \text{ yr}^{-1}$. This upper limit is higher than the expected SFR of LAEs (see e.g. Calhau et al. 2020), so, to be able to place stronger constraints on the SFR, a larger sample of LAEs, or deeper observations, are necessary.

Next, we stack the optical bands, ranging from u to y , and find no detection in the stacked full sample, with SpARCS- u , HSC- g , and HSC- r bands 3σ limiting magnitudes of ~ 27.6 , 25.1 , 24.5 respectively (using a 3" aperture), indicating that this

Table 6: Number of sources, radio flux density, RMS, and resulting SNR for LOFAR stacked images of full & optically-selected sample in NB816 and NB921.

Sample	Num. sources	$S_{150\text{MHz}}$ (μJy)	RMS (μJy)	SNR
NB816				
Full sample	226	2.4	1.7	1.4
Optically-selected sample	53	-2.6	4.0	-
Likely LAE sample	177	0.02	1.9	0.01
NB921				
Full sample	322	1.8	1.5	1.2
Optically-selected sample	49	7.4	3.9	1.9
Likely LAE sample	279	0.2	1.6	0.1

population of likely $z \sim 6$ SILVERRUSH galaxies without radio detection is consistent with $z = 6.6$ down to this magnitude limit, as expected from the initial selection method. These non-detections however do not rule out the possibility of them being low- z interlopers. We find similar results when we stack optical images for the optically-selected sample, with measurements of u , g , and r down to 3σ limiting magnitudes of 26.8, 24.2, and 23.7 respectively. In contrast, SED fitting on the full panchromatic photometry in the optically-selected sample suggested a large contamination rate (see Sect. 4.1). These sources in the optically-selected sample are typically faint in the optical and are primarily detected in the χ^2 image because of their high NIR/IR flux, indicating NIR/IR observations are critical to identify low-redshift interlopers.

Finally, we investigate the NIR and MIR data for the full LAE candidate samples (322/226 sources in NB921/NB816) by stacking the J and K images (1.25 and $2.21 \mu\text{m}$) obtained by UKIDSS, and MIR (3.6 , 4.5 , 5.8 , and $8.0 \mu\text{m}$) images obtained by Spitzer-SWIRE. We find a detection ($\text{SNR} > 3$) in all except for the $5.8 \mu\text{m}$ channel, with only a weak detection ($\text{SNR} \sim 3$) in the $8 \mu\text{m}$ channel, which again argues against strong AGN contribution.

5. Discussions

In this section we discuss the LOFAR detection rate (Sect. 5.1), the LAE sample selection, and possible biases and improvements on the selection process (Sect. 5.2).

5.1. LOFAR detection rate of SILVERRUSH LAEs

No $z > 6$ radio galaxy candidates have been found in the cross-matched LoTSS and SILVERRUSH survey. In this section, we discuss this in light of predictions made for the number densities of such sources. The expected number of radio galaxies in LoTSS in the redshift slices 5.72 ± 0.05 and 6.57 ± 0.05 , from models presented by Saxena et al. (2017), are 5.3 ± 0.4 and 1.5 ± 0.1 sources per deg^2 respectively. From these models, we would expect about 9 radio galaxies with $6.52 < z < 6.62$ to be detected in the $\sim 6 \text{ deg}^2$ overlap region between SILVERRUSH and LoTSS. However, the LAE fraction in brighter galaxies at redshifts $6 < z < 8$ is shown to be only 10-20% in the study by Schenker et al. (2012) and the distribution of $\text{Ly}\alpha$ luminosities in high- z radio galaxies ($z > 3$) is poorly constrained due to small sample sizes (see e.g. Saxena et al. 2019). Assuming this LAE fraction, only ~ 1 out of the 9 predicted sources is expected to be selected by SILVERRUSH and detected by LOFAR. Given this expectation of only 1 detection, it is not surprising that we do not find any radio emitting LAEs at $z \sim 6$. Furthermore, there are still many uncertainties in the expected number of radio detections especially at high redshift. The LoTSS Deep Field observa-

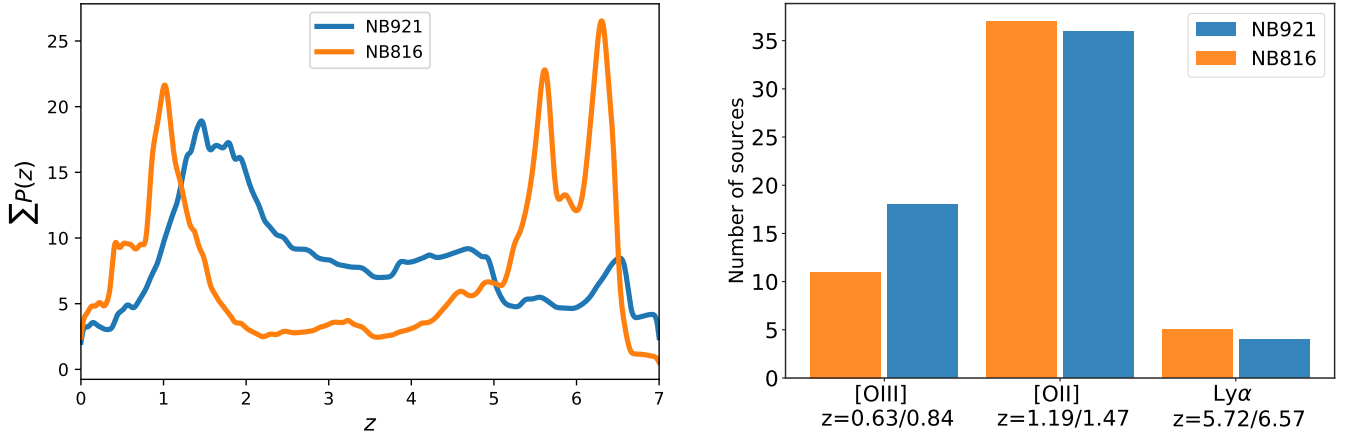


Fig. 5: Left: stacked photometric redshift probability distribution (sum of individual sources) of the optically cross-matched sample of NB921 and NB816 LAE catalogues from D20. Right: histograms of the best-fitting solutions, for SED fits performed with fixed redshift solutions corresponding to the three candidate emission lines. The known possible redshifts from the NB emission line detections and more detailed SED fitting allows to exclude other high- z solutions that are sometimes favoured by the photometric redshift code.

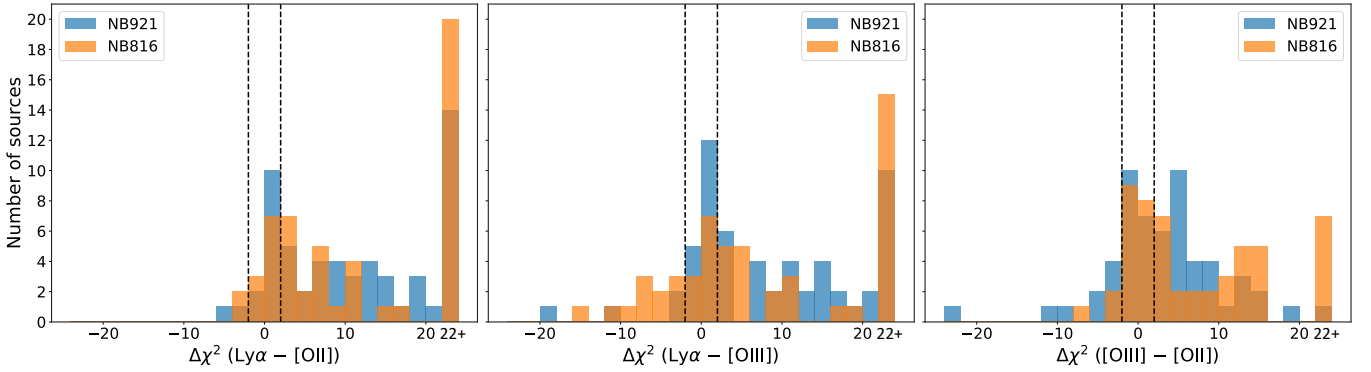


Fig. 6: Comparison of χ^2 values obtained from SED fitting models of Ly α , [OIII], and [OII] for the optical cross-matched LAE sample of NB921 and NB816. The dashed line indicates a $\Delta\chi^2$ value of 2 and -2. For higher and lower values than 2 and -2 respectively there is a preferred model, but in between the dashed lines there is no preference for either model. The bin width is set to 2.

tions probe the faint end of the radio luminosity function (RLF) due to the small volume and high sensitivity, which is poorly constrained by the models (see Saxena et al. 2017). Therefore, the uncertainties in the RLF are critical for the number density predictions in this paper. These points together provide a plausible explanation for the absence of $z > 6$ radio galaxies in the studied sample. Based on these numbers, a larger area of overlapping coverage would be necessary to find many $z > 6$ radio galaxies.

5.2. LAE selection methods

It is possible that, with the optical/IR selected samples, we are probing galaxies that are more likely to be low-redshift interlopers than the SILVERRUSH samples that do not appear in our optical/IR catalogue. Very dusty low-redshift galaxies could, for instance, have a Balmer break at 4000 Å mimicking the Lyman break (see e.g. Matthee et al. 2014, 2017). These galaxies often have red J–K colours, so constraining J–K to be flat or blue ($J - K \leq 0$) prevents contamination from these dusty low-redshift

interlopers. The J–K colours of the optically-selected samples and LOFAR detected sample are shown in Fig.9. For both the NB816 and NB921 optically-selected samples, 8 sources, out of the 53 and 58 sources respectively, do not have J and K magnitudes. Only 3 sources in NB921 satisfy $J - K \leq 0$ and these three all have high photo- z values (see Fig.9). No sources within the LOFAR-detected sample satisfy this criterion. The effect is even more apparent for NB816, where the diagonal line divides the low photo- z candidates from the high photo- z ones, except for one extremely low- z source. Therefore, NIR observations are of great value in identifying the low- z interlopers. However, note that a strict $J - K \leq 0$ cut could cause red LAEs to be rejected from the sample.

Shibuya et al. (2018a) have shown that the contamination rate depends on NB magnitude, where populations with lower magnitudes have higher contamination rates. The sources matched with our optical catalogue are brighter and likely redder on average than the full LAE sample, due to the shallower HSC observations and the source extraction from deep χ^2 image in the optical catalogue. The source extraction is based on signifi-

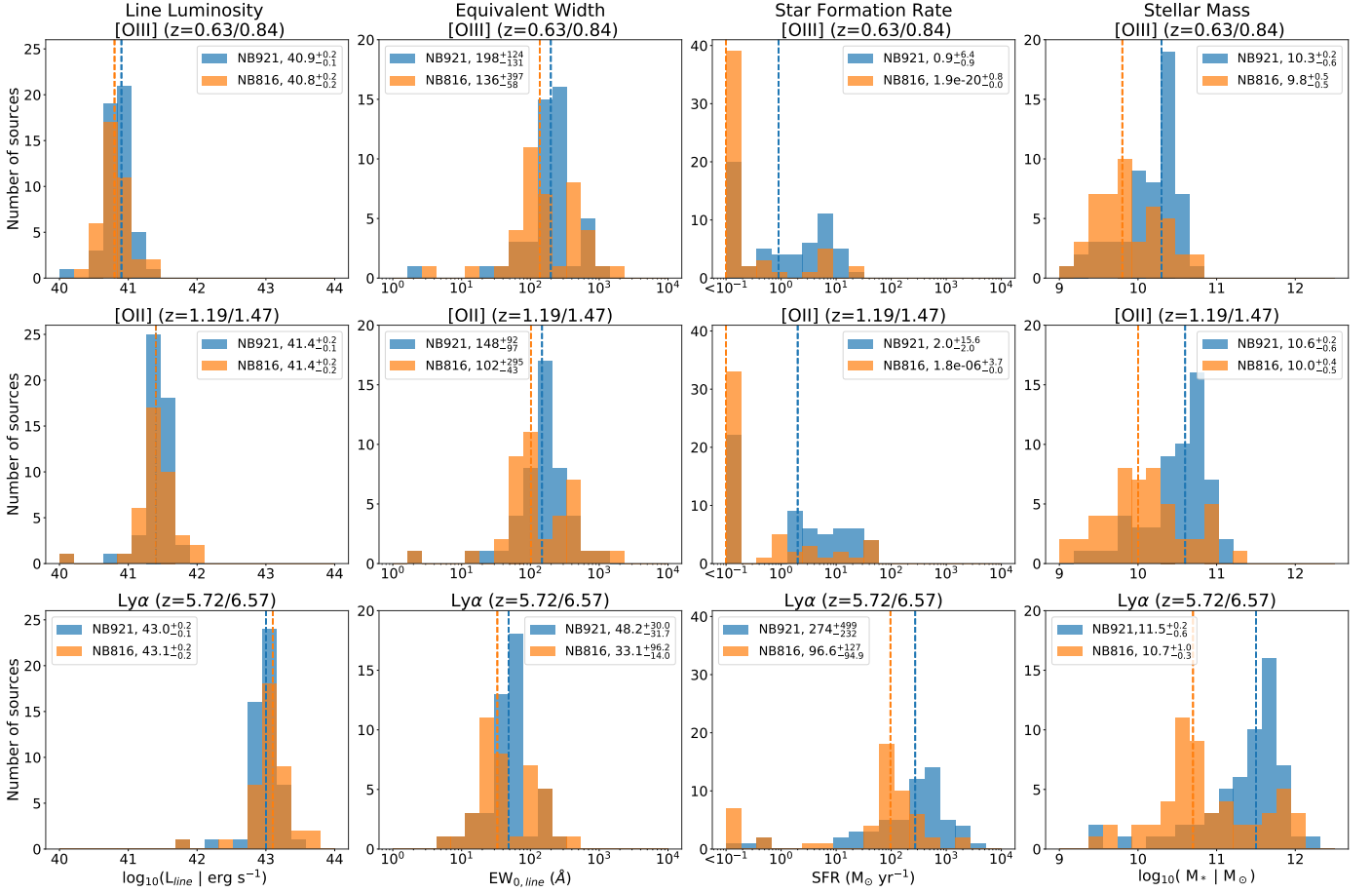


Fig. 7: Derived line luminosity, EW, SFR, and stellar mass of the optical cross-matched LAEs sample with the NB921 and NB816 LAE catalogue at different fixed redshifts of $z = 0.63/0.84$, $1.19/1.47$, $5.72/6.57$. The dashed blue and orange line indicate the median for NB921 and NB816, respectively, with the values given in the legend including errors derived from the central 68% range of the distribution. The lowest bin in SFR histograms show the sources with $\text{SFR} < 10^{-1} M_{\odot} \text{ yr}^{-1}$.

cant detection in the χ^2 stack of either optical and near-infrared or IRAC Channel 1 and 2 images. The SILVERRUSH sources that are missing from our optical sample will therefore be the ones that are fainter in the optical bands, and not red enough to be detected in the near- or mid-infrared. The χ^2 detection image did not include any y-band data, which could contribute significantly to whether or not the real LAEs are detected, since at $6 < z < 7$ sources are not detected blue-wards of the y-band where the Lyman break is situated. Therefore, the non-matched sources are more likely LAEs. If we assume that all the non-matched sources are true LAEs, then we obtain contamination rate lower limits of 23 and 20 % for the full NB816 and NB921 catalogues respectively, which would agree with the quoted 30% contamination rate from spectroscopic follow-up by SILVERRUSH (see Shibuya et al. 2018a).

To highlight the difference between the likely LAEs and the contaminants in the NB921 catalogue, we show in Fig.10 the stacked SEDs. The SED of the stacked assumed LAEs sample is consistent with $z = 6.6$, whereas the SED of the contaminants sample is consistent with $z = 1.47$. For LAEs at $z = 6.6$, the IRAC 3.6 μm filter corresponds to the [OIII]+H β emission line (see Smit et al. 2015). Since the high equivalent widths of the Ly α and [OIII]+H β lines might not be represented sufficiently well by the nebular emission line models in BAGPIPES, the NB921 filter and IRAC 3.6 μm filter have not been included

in the fits. The J–K colours are -0.02 ± 1.05 and 1.27 ± 0.11 for the stacked assumed LAEs and contaminants, respectively. This again highlights the advantage of selecting galaxies for which the NIR slope satisfies $J - K \leq 0$ when searching for high redshift LAEs. From the BAGPIPES models, we obtain stellar masses of $10^{9.6 \pm 0.2}$ and $10^{10.5 \pm 0.1} M_{\odot}$ for the SED fit of stacked LAEs and stacked contaminants respectively, and SFRs of $8.5^{+5.5}_{-3.1}$ and $< 0.4 M_{\odot} \text{ yr}^{-1}$, where the upper limit is the 84th percentile of the SFR posterior. This is consistent with the upper limit on the LAE SFR obtained from the corresponding radio stacks (see Sect.4.2).

As we concluded in Sect. 4.1, the majority of the LAE candidates that have a match in our optical catalogues are likely to be low- z interlopers of [OII] and [OIII] emitters or possibly H α at $z = 0.40$. If the sample is indeed contaminated by a large fraction of lower redshift line emitters, other NB studies might also have been affected by large fractions of contaminants, possibly leading to a misrepresentation of physical properties of LAEs.

To minimise the effects of contamination, one option for future studies is to consider the optical colour selections used to select LAE candidates. In the SILVERRUSH studies, a forced-aperture catalogue is used for statistical studies (see Shibuya et al. 2018b). This catalogue is generated with apertures at fixed positions, with a typical 3σ detection requirement that is lowered to 2σ for the g and r-bands. Colour selections are then imposed to select LAEs. The selections made are $z\text{--NB921} > 1.8$

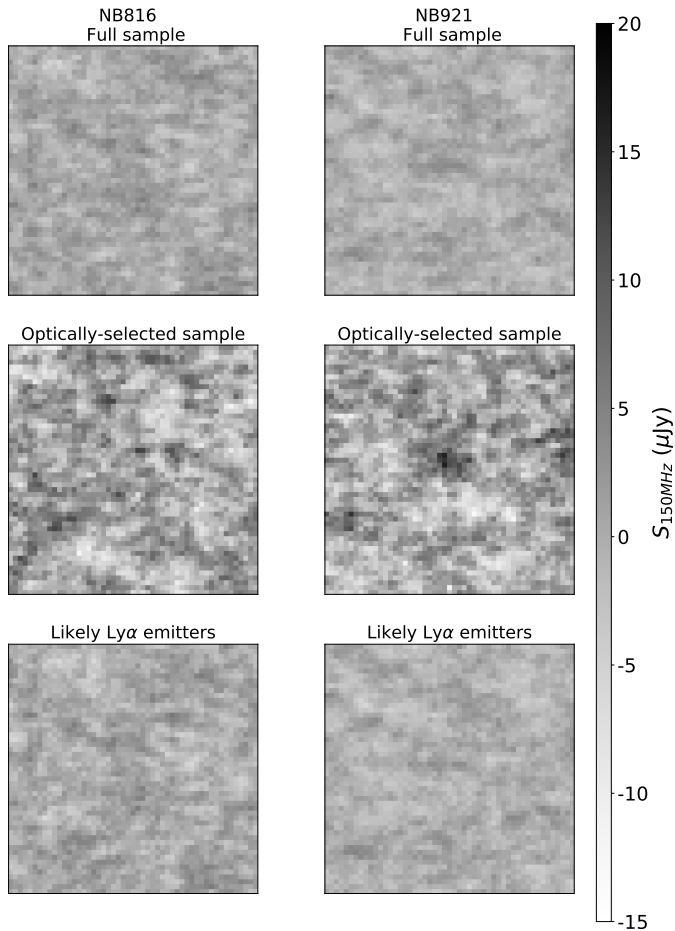


Fig. 8: Stacked LOFAR cutouts of LAE candidate samples. Left: NB816 full sample (226 sources; upper left), optically-selected sample (53 sources; middle left), and likely Ly α emitters (177 sources; bottom left). Right: NB921 full sample (322 sources; upper right), optically-selected sample (49 sources; middle right), and likely Ly α emitters (279 sources; bottom right). LOFAR detected sources and sources not overlapping with LOFAR coverage have been removed.

for $z = 6.6$ LAEs, which corresponds to $EW_{0, \text{Ly}\alpha} > 14 \text{ \AA}$ for $z \sim 6.6$ sample respectively. We now compare the numbers that would be derived from the forced SILVERRUSH catalogue, as opposed to the combined forced and unforced catalogue used in the rest of this study. The forced SILVERRUSH catalogue contains 48 sources in NB921 and 130 in NB816. The number of sources in the forced catalogues compared to the unforced catalogue is substantially lower; therefore only 2 LOFAR detected sources are found in NB921 and no sources in NB816. These two matched sources are ILTJ160227.18+544759.8 and ILTJ160930.42+544435.9, which are also likely $z = 1.5$ [OII] emitters (see top two panels in Fig. B.1). From the optical sample there are 3 and 11 in the forced catalogue of NB921 and NB816 respectively. All 3 sources in NB921 have photo- z s of $\sim 1.7 \pm 0.7$, suggesting these are low-redshift interlopers. However, of the 11 sources in NB816 forced catalogue, 10 have a photo- z in the range $4.5\text{--}5.8 \pm 1.5$ (the one remaining source has a photo- z of 0.5 ± 0.9), suggesting that the LAE fraction is likely high. This implies that, at least in the case of $z = 5.7$ LAEs, the contamination rate could be significantly lowered by making stricter cuts.

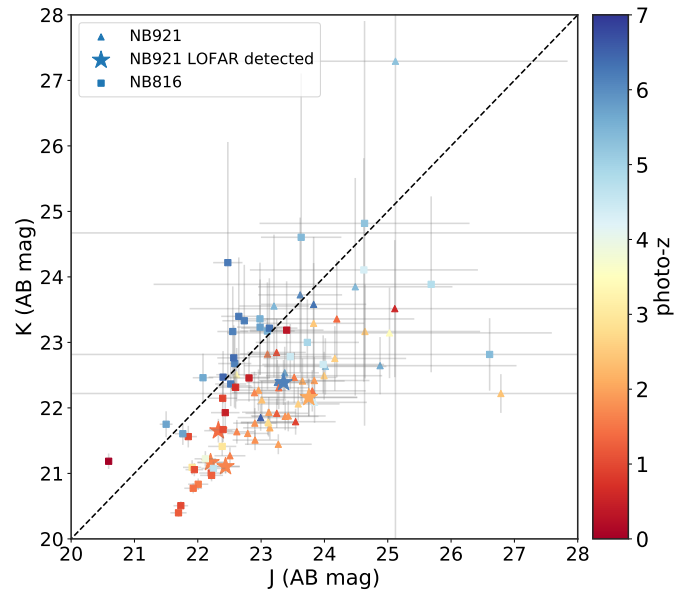


Fig. 9: UKIDSS J and K band magnitude of the optically-selected sample and LOFAR detected sample, with colour coded photo- z s from D20. The diagonal (dashed) line indicates the $J-K \leq 0$ criterion. The sources above this line satisfy this criterion and also have notably higher photo- z values.

Many LAE narrowband surveys apply similar cuts for selection, but differ in strictness on the cuts and availability of multiple bands (e.g. Ouchi et al. 2008; Santos et al. 2016; Matthee et al. 2015). Other studies have also shown that the majority of LAEs have $EW_0 > 50 \text{ \AA}$ (e.g. Ouchi et al. 2008). However, Sobral et al. (2018b) also show by using a large compilation of spectroscopic redshifts that, even though their sample of high EW candidate line-emitters is dominated by likely LAEs, there is still a significant population of H α , [OIII]+H β , and [OII] emitters. Since the observed EW_0 distribution in these lines is similar to Ly α (see e.g. Khostovan et al. 2016; Hayashi et al. 2018), it can be challenging to clean the samples further using just NB and BB optical surveys. The contamination rates we obtain are comparable to those found by Matthee et al. (2014), who find a 90% contamination in LAE samples selected on the Lyman break criteria. However, their survey is targeting higher redshifts of $z > 7$, so contamination is expected to be higher.

This work further highlights that ancillary infrared observations are critical for the LAE sample selection. Although the J to $4.5 \mu\text{m}$ data in EN1 is significantly shallower than available in smaller fields such as COSMOS, it can still play an important role in minimising interlopers in wide field LAE samples. Decreasing the contamination rate is critical for improving the constraints on the Ly α luminosity function and clustering measurements, and for reducing the time and expense required for spectroscopic follow-up observations.

6. Summary

In this paper, we have studied the LOFAR and optical/IR properties of SILVERRUSH Ly α emitters (LAEs) in EN1 using the latest LoTSS Deep Fields observations and cross-matched multi-wavelength catalogue ranging from UV to FIR. The publicly available SILVERRUSH catalogues of EN1 consist of 229 and

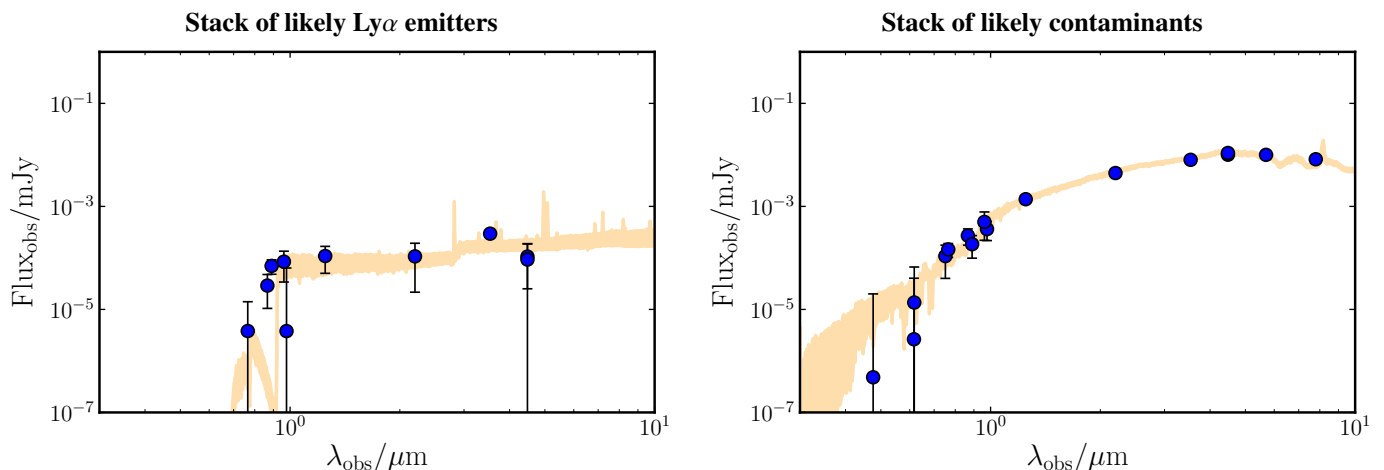


Fig. 10: Stacked SEDs (blue) with posterior median SED models (yellow). Left: stacked likely $z = 6.6$ LAE candidates in the NB921 catalogue, consisting of the non-matched sample and probable LAEs in the matched sample. Right: stacked likely contaminants in the NB921 matched optical sample, fitted with fixed $z = 1.47$. The stack of likely contaminants shows brighter NIR and MIR fluxes highlighting the importance of using NIR data in LAE selection.

349 LAE candidates at $z = 5.7$ and $z = 6.6$, respectively, identified with the Hyper Suprime Cam NB816 and NB921 narrowbands. We found 5 LOFAR detected LAEs in the NB921 SILVERRUSH catalogue. These 5 sources have LOFAR radio fluxes in the range of 0.09–0.35 mJy. We perform SED fitting on the radio detected sources using BAGPIPES and find that all 5 sources are likely $z = 1.5$ [OII] emitters, based on χ^2 analysis and derived physical properties. We therefore do not find any promising $z > 6$ radio galaxy candidates and cannot put any constraints on the AGN fraction of LAE at high redshift.

In light of the high contamination rate of radio-selected SILVERRUSH sources, we perform the same SED fitting procedure on the wider LAE population, by cross-matching the SILVERRUSH NB816 and NB921 catalogues with the LoTSS Deep Field multi-wavelength catalogue. We find 53/58 sources with matches in the LoTSS multi-wavelength catalogue, out of the 229/349 SILVERRUSH LAE candidates. The resulting $\Delta\chi^2$ values suggest very high contamination rates of 81(92) and 81(91)% for the NB816 and NB921 catalogue sample, respectively, where the values in the parenthesis include the sources which are also likely contaminants but not statistically significant. This is consistent with the photo- z s from D20 given the large photo- z uncertainties at these high redshifts and faint magnitudes. Of all the contaminated sources, $63 \pm 31\%$ are likely to be [OII] and $6 \pm 31\%$ [OIII] emitters for NB816, and similarly $55 \pm 28\%$ are likely to be [OII] and $17 \pm 28\%$ [OIII] emitters for NB921 (with the remainder inconclusive). However, the non-matched sources are fainter and bluer and therefore more likely to be true LAEs, which would be in line with the 30% contamination rate obtained from spectroscopic follow-up by SILVERRUSH (Shibuya et al. 2018a).

Most importantly, in agreement with Matthee et al. (2014, 2017), we find that a $J-K \leq 0$ cut can identify a large fraction of red and dusty low- z interlopers, which are not removed by the standard LAE criteria due to the low optical SNRs and the Balmer break mimicking the Lyman break. These results highlight the need for ancillary infrared observations in LAE sample selection to minimise the amount of interlopers.

We stack radio image cutouts in the positions of all LAE candidates and find no significant detection with LOFAR when removing the radio detected sources. Removing the likely con-

taminating sources and stacking the LOFAR data for the 4 and 5 remaining matched sources in NB816 and NB921 together with the 173 and 274 non-matched sources, yields a 2σ upper limit on the SFR of ~ 53 and $56 M_{\odot} \text{ yr}^{-1}$ respectively. In the optical stacked u, g, r measurements, we do not find a detection down to $\sim 27.6, 25.1, 24.5$ magnitude respectively, which is consistent with the Lyman break at $z = 6.6$. The stacked IRAC SWIRE 3.6 and $4.5 \mu\text{m}$ and UKIDSS J and K images do yield a detection ($\text{SNR} > 3$), but there is no significant detection ($\text{SNR} \sim 2$) in the IRAC SWIRE 5.8 and $8.0 \mu\text{m}$ measurements.

The current 180 hours of observations in EN1 are expected to reach 500 hours within the next two years. This will result in factor of ~ 2 deeper observations down to $11 \mu\text{Jy beam}^{-1}$, enabling the study of the fainter and higher redshift radio population. Furthermore, the WEAVE-LOFAR spectroscopic survey (Smith et al. 2016; commissioning early 2021) will obtain around 10^6 optical spectra of LOFAR detected radio sources, allowing accurate redshift determination and source classification. This survey will provide the opportunity to reveal high redshift ($z > 6$) radio galaxies, which are necessary to advance our current understanding of galaxy formation and evolution and to use as probes of the epoch of reionization. Follow-up observations with existing 10m class optical telescopes and up-coming facilities, such as *James Webb Space Telescope* and *European Extremely Large Telescope*, will enable detailed studies of these high redshift galaxies for the first time.

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Table A.1: Overview of parameters used in this work for fitting to our data in BAGPIPES.

Symbol	Prior value range	Parameter description
<i>SFH</i>		
τ	[0,15] Gyr	peak time of SFR
α	[0.01,1000]	rising power slope
β	[0.01,1000]	falling power slope
$M_{*,\text{formed}}$	[7,14] $\log_{10}(M_{\odot})$	amount of stellar mass formed
metallicity	[0,2.5] M_{\odot}	metallicity of galaxy
<i>Dust</i>		
type	Calzetti	attenuation law
A_V	[0,6]	absolute attenuation in V band
umin	[0.1,25]	lower limit of starlight intensity distribution
γ	[0.0001,1.0]	fraction of stars at umin
qpah	[0.1,4.5]	PAH mass fraction
<i>Nebular</i>		
$\log_{10}(U)$	-3	ionisation parameter

Appendix A: BAGPIPES model parameters

Table A.1 shows the parameters we use to fit to our data in BAGPIPES (Carnall et al. 2018). The star formation history (SFH) is modelled using a double-power-law given by

$$SFR(t) \propto \left[\left(\frac{t}{\tau} \right)^{\alpha} + \left(\frac{t}{\tau} \right)^{-\beta} \right]^{-1} \quad (\text{A.1})$$

with falling power slope α , rising power slope β , and τ the time of peak star formation. A prior distribution uniform in \log_{10} is used for α and β . For further information on the model components visit https://bagpipes.readthedocs.io/en/latest/model_components.html.

Appendix B: SED fits of LOFAR detected sources

The SED fits and reduced χ^2 as a function of redshift are given in Figure B.1 for 4 sources in the LOFAR detected sample.

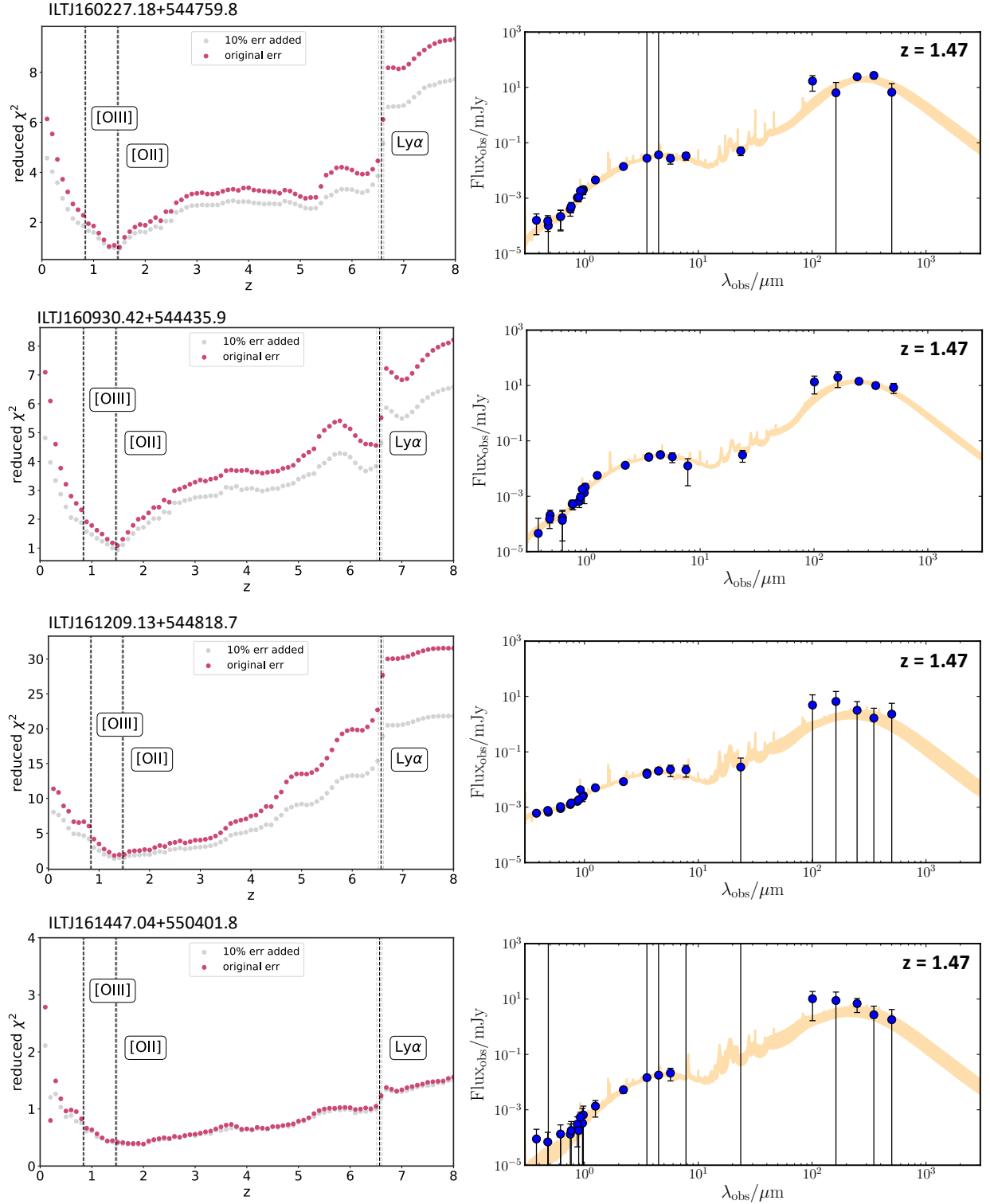


Fig. B.1: SED fitting results obtained using BAGPIPES for ILTJ160227.18+544759.8, ILTJ160930.42+544435.9, ILTJ161209.13+544818.7, and ILTJ161447.04+550401.8 (in order from top to bottom). Left: reduced χ^2 as function of redshift. The redshifts corresponding to [OIII], [OII], and Ly α emission lines are indicated and both resulting reduced χ^2 values using normal error (pink) and 10% extra flux error (grey) are shown. Right: posterior median model obtained from BAGPIPES (yellow) with $z = 1.47$ and over-plotted photometric fluxes as function of wavelength (blue).